Friedman (FLRW) universe

The *Friedman universe* is a cosmological model where the universe is assumed to be homogeneous, isotropic, and filled with *ideal fluid*. The assumption of homogeneity and isotropicity of the universe is often referred to as the *cosmological principle*. Empirically, it seems to by justified on scales larger that 100 Mpc. The isotropic and homogeneous model is sometimes called the *Standard Model* of present-day cosmology. It is most often refered to as Friedman-Lemaitre-Robertson-Walker model (or FLRW model, for short).

The *Friedman equation* is the Einstein equation applied to the Friedman universe as a whole. It describes the temporal evolution of a Friedman universe.

Spaces with constant curvature

A homogeneous and isotropic universe is a *space with constant curvature*.

Two-dimensional spaces of constant curvature are three-dimensional spheres (positive curvature), pseudo-spheres (negative curvature), and planes (zero curvature).

On a sphere the length element in the ordinary spherical coordinates is given as

$$dl^2 = a^2\left(d\theta^2 + \sin^2\theta d\phi^2\right) \ , \tag{1}$$

where a is the radius of the sphere.

Let us introduce the polar coordinates $\{r, \phi\}$ on the sphere, with r measuring the distance to the north pole. The length of a circle around north pole, $\theta = \text{const}$, is equal $2\pi a \sin\theta$. Therefore if we want the circumference of the circle to be equal $2\pi r$, we need to define $r = a\sin\theta$.

The length element in the polar coordinates becomes

$$dl^2 = \frac{dr^2}{1 - \frac{r^2}{a^2}} + r^2 d\phi^2 \,. \tag{2}$$

On a pseudo-sphere, where the curvature is negative, the length element is correspondingly

$$dl^2 = \frac{dr^2}{1 + \frac{r^2}{a^2}} + r^2 d\phi^2 \,. \tag{3}$$

In angular coordinates $r = a\sinh\theta$ the latter becomes

$$dl^2 = a^2\left(d\theta^2 + \sinh^2\theta d\phi^2\right) \ . \tag{4}$$

On a plane, where the curvature is zero, the length element is

$$dl^2 = dr^2 + r^2 d\phi^2 \,. \tag{5}$$

Analogously, a three-dimensional space with constant curvature can have one of the following three possible geometries:

- *flat* (zero curvature),

$$dl^2 = dr^2 + r^2(d\theta^2 + \sin^2\theta d\phi^2) = a^2 d\Omega_3^2 \,, \tag{6}$$

where

$$d\Omega_3^2 = d\chi^2 + \chi^2(d\theta^2 + \sin^2\theta d\phi^2) \,, \tag{7}$$

where $r = a\chi$, $\chi \in [0, \infty]$;

- *closed* (positive curvature),

$$dl^2 = \frac{dr^2}{1 - \frac{r^2}{a^2}} + r^2(d\theta^2 + \sin^2\theta d\phi^2) = a^2 d\Omega_3^2 \,, \tag{8}$$

where

$$d\Omega_3^2 = d\chi^2 + \sin^2\chi(d\theta^2 + \sin^2\theta d\phi^2), \tag{9}$$

where $r = a\sin\chi$, $\chi \in [0,\pi]$;

- and *open* (negative curvature),

$$dl^2 = \frac{dr^2}{1+\frac{r^2}{a^2}} + r^2(d\theta^2 + \sin^2\theta d\phi^2) = a^2 d\Omega_3^2\,, \tag{10}$$

where

$$d\Omega_3^2 = d\chi^2 + \sinh^2\chi(d\theta^2 + \sin^2\theta d\phi^2), \tag{11}$$

where $r = a\sinh\chi$, $\chi \in [0,\infty[$.

Friedman equation

Friedman metric[1] is a (generic) synchronous metric in a homogeneous and isotropic universe,

$$ds^2 = dt^2 - a(t)^2 d\Omega_3^2\,, \tag{12}$$

where $a(t)^2 d\Omega_3^2$ is the line element in a three-dimensional space of constant curvature (open, closed, or flat), and $a(t)$ is the time-dependent *scale factor* of the universe.

Closed universe In a closed Friedman universe the metric is

$$ds^2 = a^2\left(d\eta^2 - d\chi^2 - \sin^2\chi(d\theta^2 + \sin^2\theta d\phi^2)\right)\,, \tag{13}$$

where $r = a\sin\chi$, and η is the scaled time coordinate,

$$dt = a d\eta\,. \tag{14}$$

Now the following Maxima script

```
derivabbrev:true; /* a better notation for derivatives */
load(ctensor); /* load the package to deal with tensors */
ct_coords:[eta,chi,theta,phi]; /* our coordiantes: eta chi theta phi */
depends([a],[eta]); /* the scale, a, depends on eta */
lg:ident(4); /* set up a 4x4 identity matrix */
lg[1,1]: a^2; /* g_{\eta\eta} = a^2 */
lg[2,2]:-a^2; /* g_{\chi\chi} = -a^2 */
lg[3,3]:-a^2*sin(chi)^2; /* g_{\chi\chi} */
lg[4,4]:-a^2*sin(chi)^2*sin(theta)^2; /* g_{\phi\phi} */
cmetric(); /* compute the prerequisites for further calculations */
christof(mcs); /* print out the Christoffel symbols, mcs_{bca}=\Gamma^a_{bc} */
uricci(true); /* print out the elements of the Ricci tensor */
scurvature(); /* Ricci scalar curvature */
```

calculates the components of the Ricci tensor[2] and the Ricci scalar for this metric,

$$R^\eta_\eta = \frac{3}{a^4}(a'^2 - aa'')\,,\; R^\chi_\chi = R^\theta_\theta = R^\phi_\phi = -\frac{1}{a^4}(2a^2 + a'^2 + aa'')\,,\; R = -\frac{6}{a^3}(a + a'')\,, \tag{15}$$

where prime denotes the η-derivative.

[1]also referred to as Friedman-Lemaitre-Robertson-Walker in different combinations.

[2]The script also calculates the Christoffel symbols, $\Gamma^\eta_{\eta\eta} = \Gamma^\chi_{\eta\chi} = \Gamma^\theta_{\eta\theta} = \Gamma^\phi_{\eta\phi} = \Gamma^\eta_{\chi\chi} = \frac{a'}{a}$, $\Gamma^\theta_{\chi\theta} = \Gamma^\phi_{\chi\phi} = \cot\chi$, $\Gamma^\eta_{\theta\theta} = \frac{a'}{a}\sin^2\chi$, $\Gamma^\chi_{\theta\theta} = -\cos\chi\sin\chi$, $\Gamma^\phi_{\theta\phi} = \cot\theta$, $\Gamma^\eta_{\phi\phi} = \frac{a'}{a}\sin^2\chi\sin^2\theta$, $\Gamma^\chi_{\phi\phi} = -\cos\chi\sin\chi\sin^2\theta$, $\Gamma^\theta_{\phi\phi} = -\cos\theta\sin\theta$.

The energy-momentum tensor for the perfect fluid is[3]

$$T_{ab} = (\epsilon + p)u_a u_b - p g_{ab}\,, \tag{16}$$

where ϵ is the rest-energy density and p is the pressure. In synchronous Friedman coordinates the matter is at rest and the 4-velocity is $u^b = \{\frac{1}{a}, 0, 0, 0\}$.

Correspondingly, the Einstein equation

$$R^a_b - \frac{1}{2}R\delta^a_b = \kappa T^a_b \tag{17}$$

then has the ${}^\eta_\eta$ component

$$\frac{3}{a^4}(a^2 + a'^2) = \kappa\epsilon, \tag{18}$$

and the three identical spatial components,

$$\frac{1}{a^4}(a^2 + 2aa'' - a'^2) = -\kappa p\,, \tag{19}$$

called the *Friedman equations* for a closed universe,

If the relation between ϵ and p, called the *equation of state* of the matter, is known, the energy density ϵ can be determined as function of a from the energy conservation equation. The latter must have the form

$$dE = -pdV\ , \tag{20}$$

where V is a volume element in the Friedman universe, and $E = \epsilon V$ is the energy content of this volume. Since the volume is proportional to a^3, and both ϵ and a in a Friedman universe can only depend on time, equation (20) can be rewritten as

$$(\epsilon a^3)' + p(a^3)' = 0\ . \tag{21}$$

It is easy to show, that equation (21) actually follows from the Friedman equations (18) and (19).

The energy conservation equation (20) can also be written as

$$\frac{3da}{a} = -\frac{d\epsilon}{\epsilon + p}\ . \tag{22}$$

When the dependence $\epsilon(a)$ is found by integration of the energy conservation equation (22), the solution to the Friedman equation can be obtained as the integral

$$\eta = \pm\int \frac{da}{a\sqrt{\frac{1}{3}\kappa\epsilon a^2 - 1}}\,. \tag{23}$$

Open universe In an open Friedman universe the metric is

$$ds^2 = a^2\left(d\eta^2 - d\chi^2 - \sinh^2\chi(d\theta^2 + \sin^2\theta d\phi^2)\right)\ , \tag{24}$$

where $r = a\sinh\chi$, and $ad\eta = dt$. This metric can be obtained from the closed universe metric (13) by a formal substitution

$$\{a, \eta, \chi\} \to \{ia, i\eta, i\chi\}\,. \tag{25}$$

[3]Interpreting the equations $T^{ab}{}_{;b} = 0$ as conservation laws leads to the following interpretations of the components of the energy-momentum tensor: T^{00} is energy density, $T^{0\alpha}$ is momentum density, $T^{\alpha\alpha}$ is pressure, and $T^{\alpha\beta}$ wher $\alpha \neq \beta$ is the shear stress. For a perfect fluid at rest the shear stress is zero and the momentum is also zero. Thus in the frame where the element of the liquid is at rest, $u^a = \{1, 0, 0, 0\}$, the energy-momentum tensor is diagonal with components ϵ, p, p, p where ϵ is the rest-energy and p is the pressure. Apparently, the covariant form must then be $T^{ab} = (\epsilon + p)u^a u^b - pg^{ab}$.

Therefore the Friedman equation for an open universe can be readily obtained from (18) by the substitution (25),

$$\frac{3}{a^4}(a'^2 - a^2) = \kappa\epsilon, \tag{26}$$

with the integral solution,

$$\eta = \pm \int \frac{da}{a\sqrt{\frac{1}{3}\kappa\epsilon a^2 + 1}} \,. \tag{27}$$

Exercises

1. Calculate the (diagonal components of the) Ricci tensor in Friedman coordinates for a closed universe.
2. Argue that the energy conservation equation,
$$(\epsilon a^3)' + p(a^3)' = 0\,, \tag{28}$$
follows from the field equations
$$\frac{3}{a^4}(a^2 + a'^2) = \kappa\epsilon\,, \tag{29}$$
and
$$\frac{1}{a^4}(a^2 + 2aa'' - a'^2) = -\kappa p\,. \tag{30}$$
3. Consider a flat (Euclidean) isotropic universe with the metric
$$ds^2 = dt^2 - a^2(t)(dx^2 + dy^2 + dz^2)\,.$$
and investigate its temporal development for matter and radiation dominated universes.
   (a) Calculate the Christoffel symbols,
   $$\Gamma^x_{tx} = \Gamma^y_{ty} = \Gamma^z_{tz} = \frac{\dot{a}}{a}\,,\ \Gamma^t_{xx} = \Gamma^t_{yy} = \Gamma^t_{zz} = a\dot{a}\,.$$
   (b) Calculate the Ricci tensor and the Ricci scalar,
   $$R^t_t = -3\frac{\ddot{a}}{a}\,,\ R^x_x = R^y_y = R^z_z = -\frac{\ddot{a}}{a} - 2\frac{\dot{a}^2}{a^2}\,.$$
   (c) Write down the ${}^t_t$ component of the Einstein equation with perfect fluid,
   $$3\frac{\dot{a}^2}{a^2} = \kappa\epsilon\,.$$
   (d) Write down the energy conservation equation,
   $$\frac{dV}{V} = -\frac{d\epsilon}{(\epsilon + p)} \Rightarrow 3\ln(a) = -\int \frac{d\epsilon}{(\epsilon + p)}\,.$$
   (e) Integrate the equations for a matter dominated universe ($p = 0$, $\epsilon = \mu$),
   $$\mu a^3 = \text{const}\,,\ a =\propto t^{2/3}\,.$$
   (f) Integrate the equations for a radiation dominated universe ($p = \frac{\epsilon}{3}$),
   $$\epsilon a^4 = \text{const}, a =\propto t^{1/2}$$
4. Calculate the volumes of the closed and open universes.