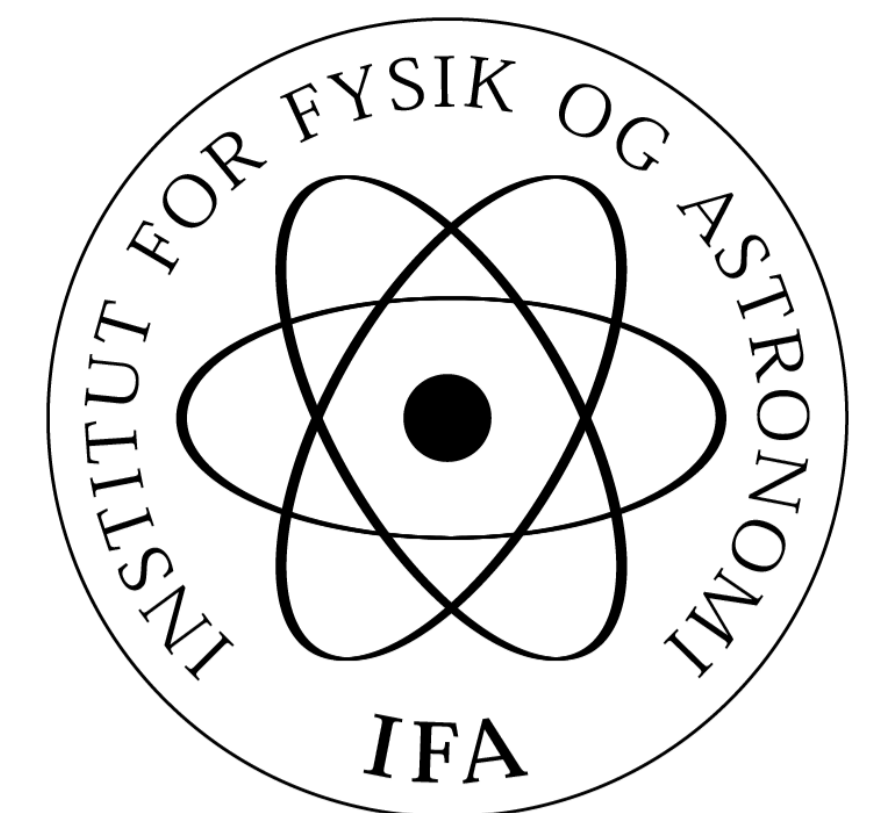


# Laboratory astrophysics at FLASH: XUV photofragmentation of the water cation $H_2O^+$

Max-Planck-Institut  
für Kernphysik

H. B. Pedersen<sup>1</sup>, C. Domesle<sup>2</sup>, L. Lammich<sup>1</sup>, S. Dziarczyński<sup>4</sup>, N. Guerassimova<sup>4</sup>,  
L. S. Harbo<sup>1</sup>, O. Heber<sup>3</sup>, B. Jordon-Thaden<sup>2</sup>, R. Treusch<sup>4</sup>,  
T. Arion<sup>5</sup>, M. Förstel<sup>5</sup>, U. Hergenhan<sup>5</sup>, M. Stier<sup>5</sup>, and A. Wolf<sup>2</sup>



<sup>1</sup>Department of Physics and Astronomy, Aarhus University, Denmark, <sup>2</sup>Max-Planck-Institut für Kernphysik, Heidelberg, Germany, <sup>3</sup>Department of Particle Physics, Weizmann Institute of Science, Rehovot, Israel, <sup>4</sup>HASYLAB, DESY at Hamburg, Germany, <sup>5</sup>Max-Planck-Institut für Plasmaphysik, Garching, Germany

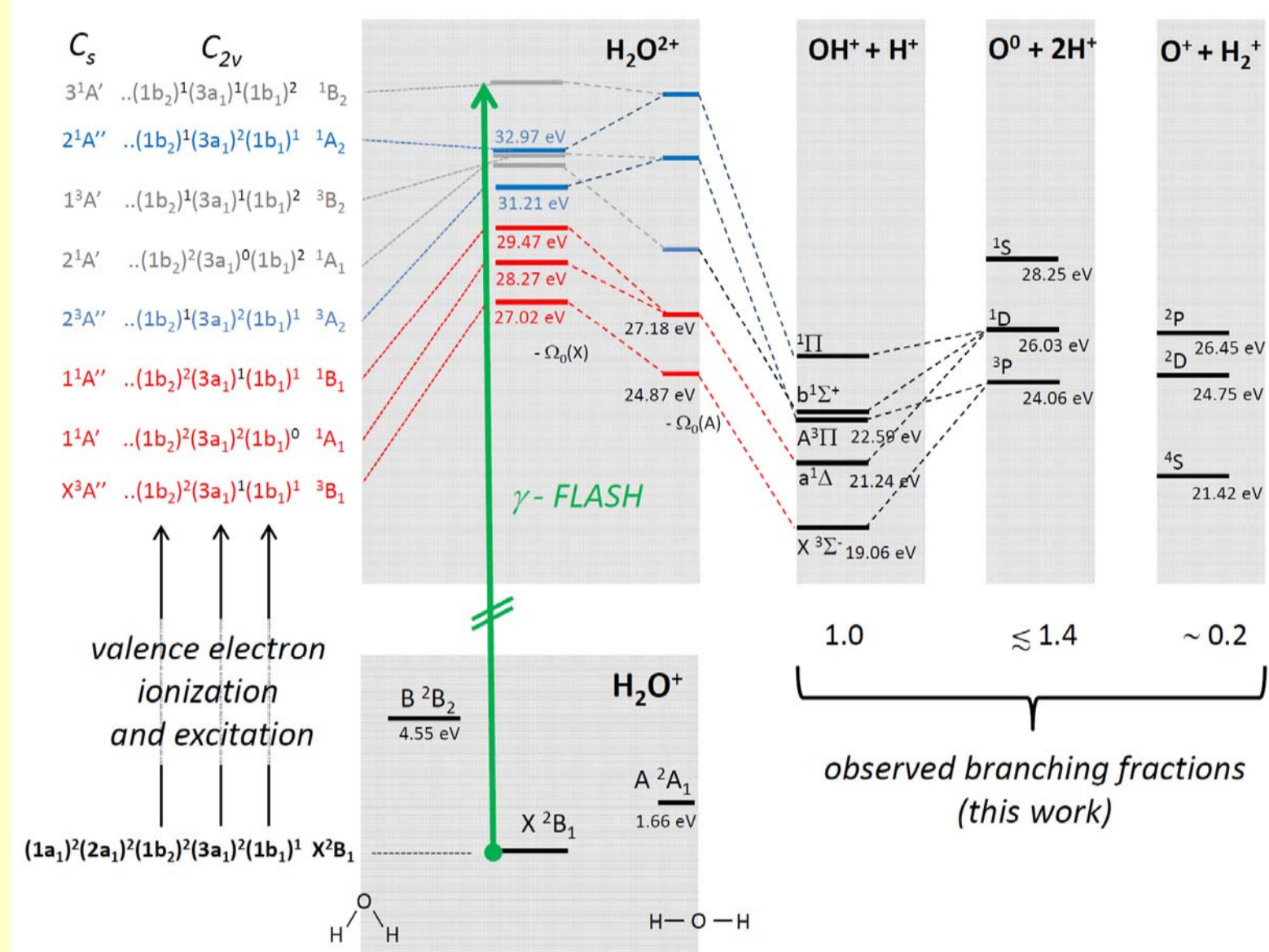
## XUV photofragmentation of $H_2O^+$

The water cation  $H_2O^+$  occurs in several natural environments where molecular gas is exposed to ionizing radiation. It has been directly observed in the upper atmosphere [1] and in the tails of several comets [2–7] through its visible emission spectrum. These observations have been used to indirectly identify the presence of neutral water in comets, since  $H_2O$  molecules themselves do not show a visible emission spectrum. Recently, the astrophysical importance of  $H_2O^+$  has been highlighted by several observations of both  $H_2O^+$  as well as other oxygen hydride ions ( $OH^+$  and  $H_3O^+$ ) using the Herschel Space Observatory [8–11].

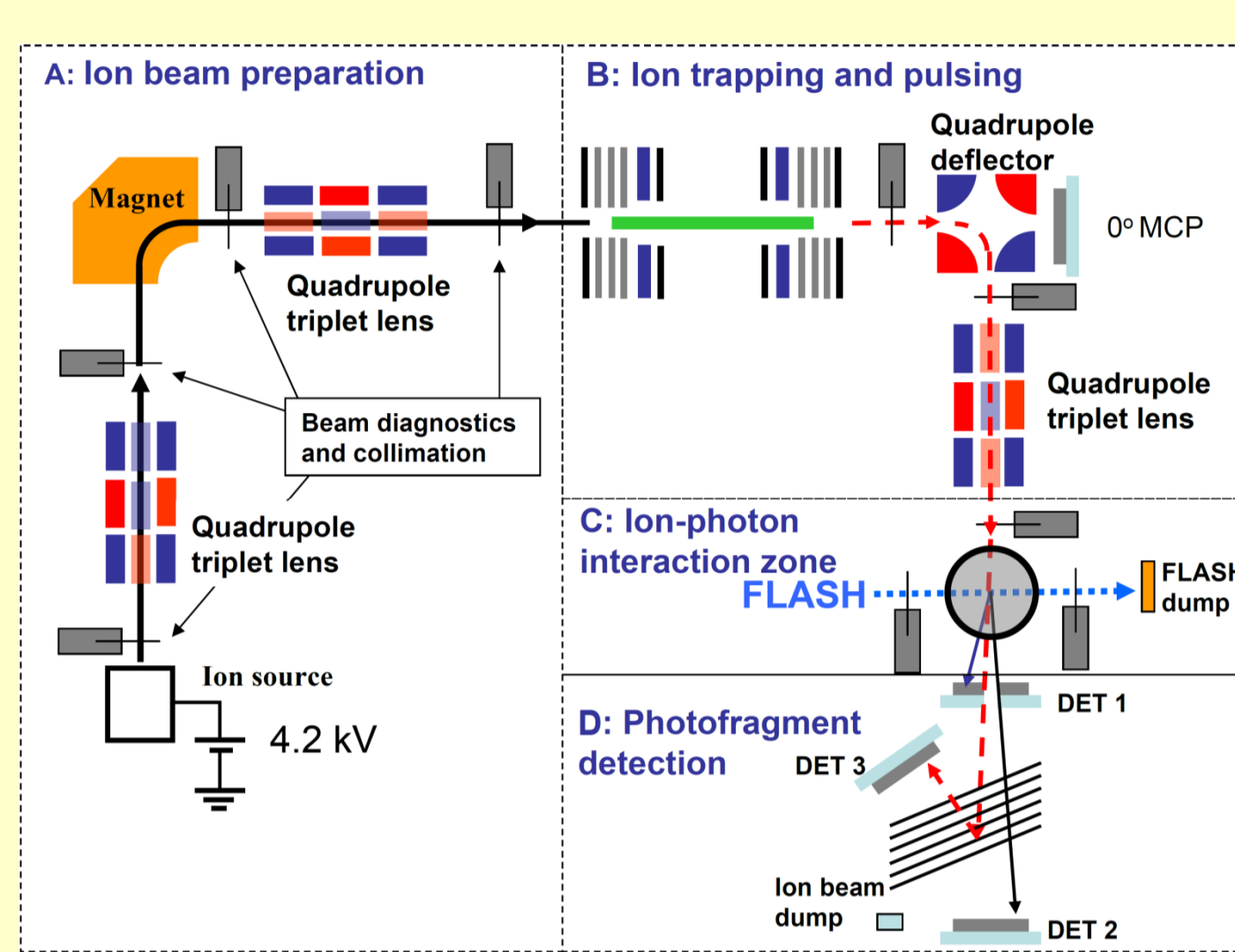
Reactions of oxygen and its smaller hydrides are part of models of the interstellar chemistry in various environments [12–15], where in particular  $H_2O^+$  ions are formed in reaction of  $H_2$  and  $OH^+$ , while being depleted by further reactions with  $H_2$  to form  $H_3O^+$  and by dissociative recombination. The fragmentation of water cations under ionizing radiation is not considered in present models.

Using FLASH in combination with techniques based on accelerated ion beams we have investigated the fragmentation of isolated  $H_2O^+$  ions under ionizing radiation in the form of mono-energetic photons for two wavelengths, 35.0 nm and 21.8 nm, in the extreme ultraviolet (XUV) regime.

In a molecular physics perspective, we here make the first investigation of the potential energy surfaces of the water dication starting from the mono-cation [16].



## TIFF experiment with new analyzing system for charged photofragments



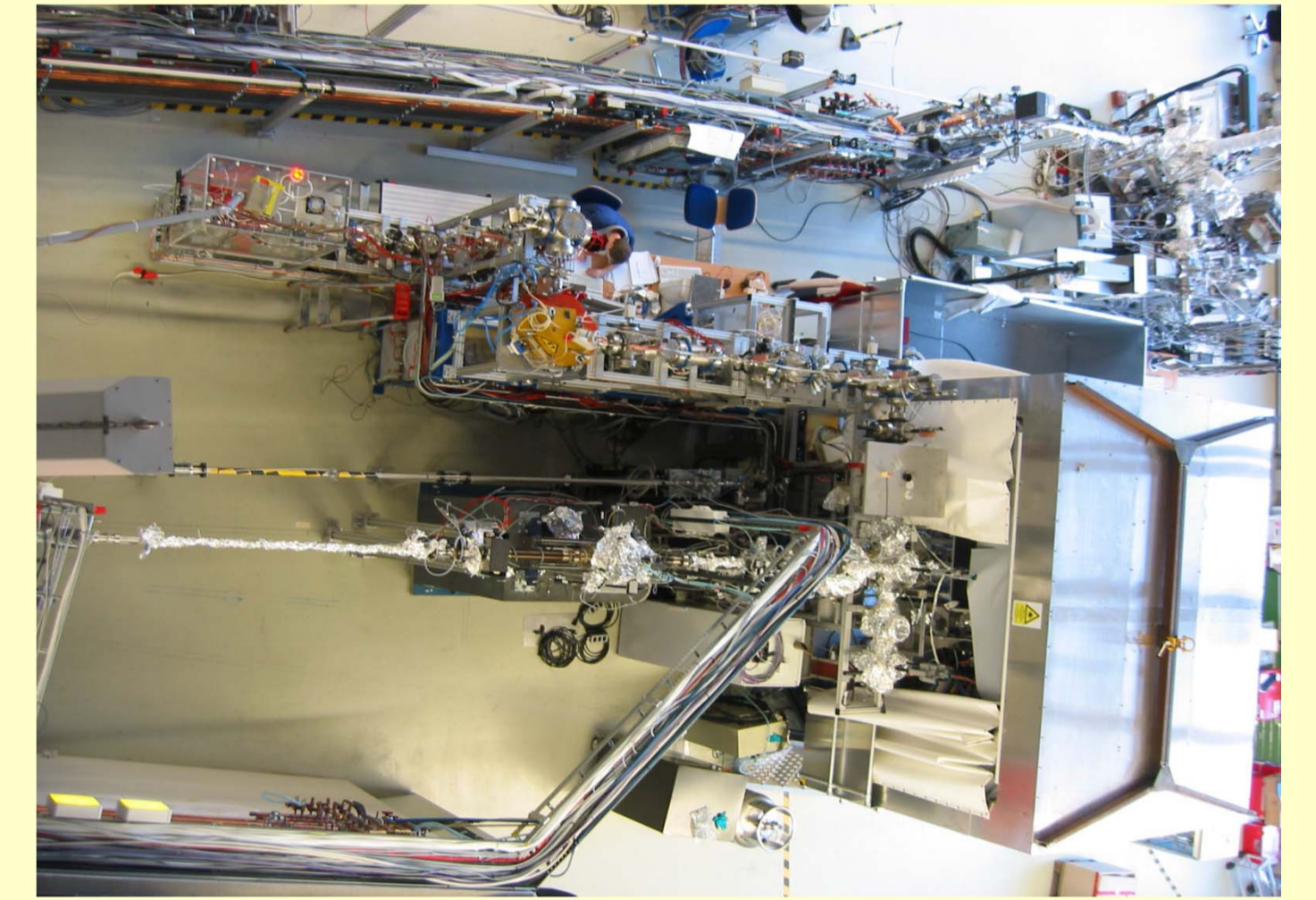
TIFF

Trapped

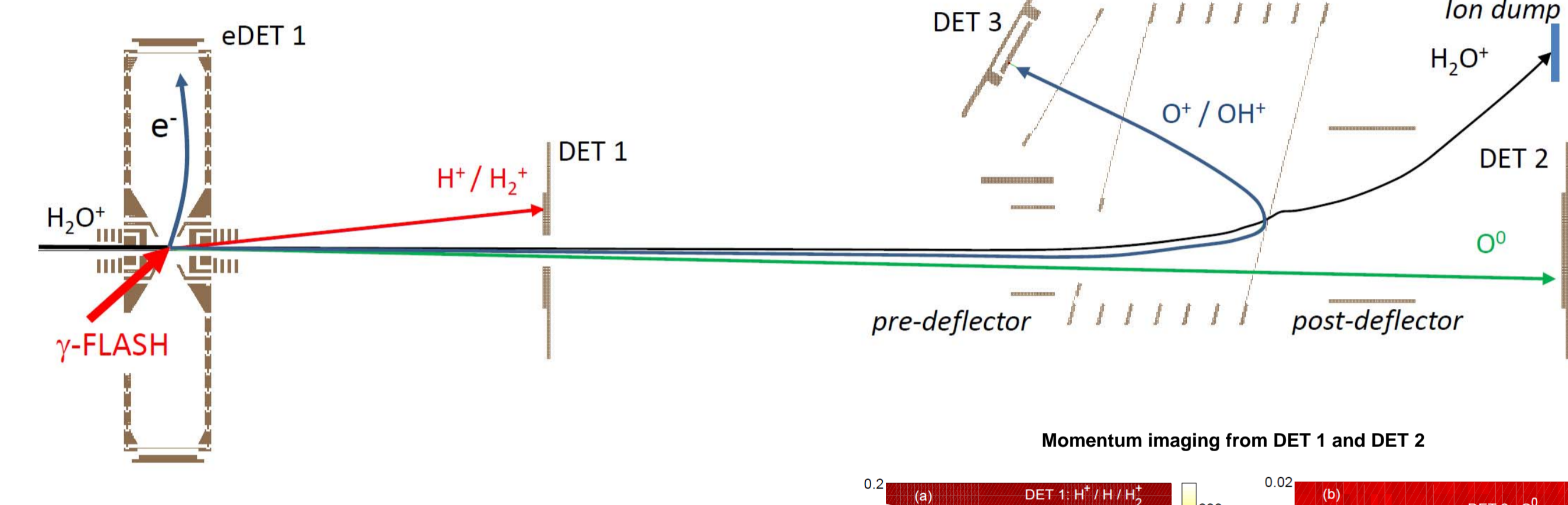
Ion

Fragmentation at

FLASH



### Photofragment detection



### FLASH

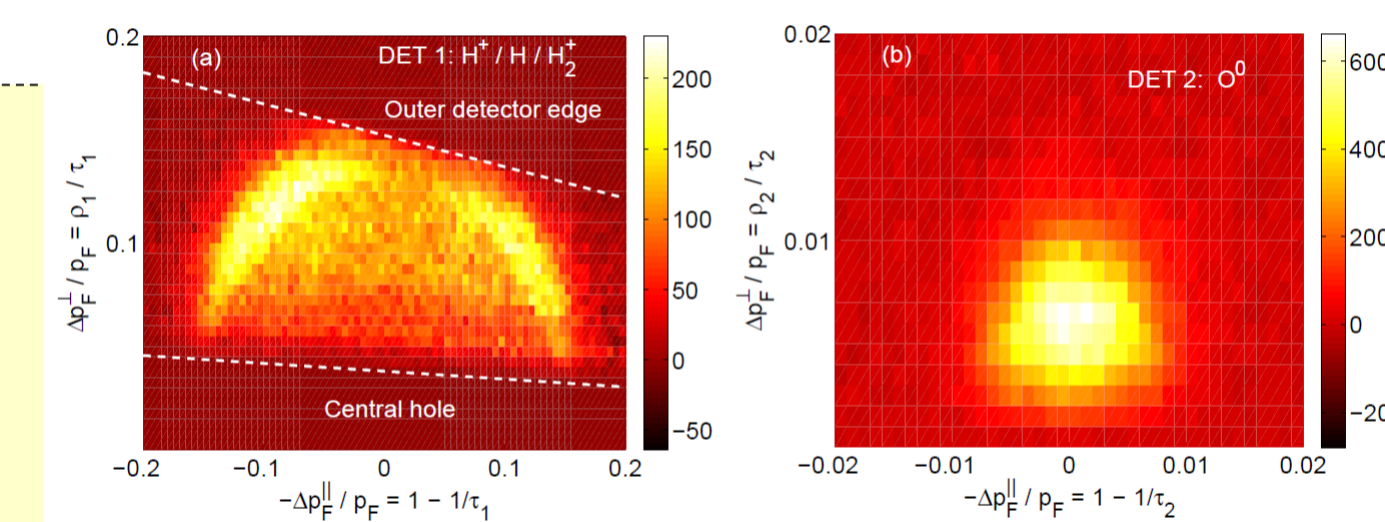
- $\lambda = 21.8 \text{ nm} / 35.0 \text{ nm}$
- $E = 30\text{-}40 \mu\text{J/pulse}$
- 50 pulses at 200kHz (5 $\mu\text{s}$ )

### TIFF

- $E_i = 4.2 \text{ keV}$
- $I = 20 \text{ nA}$
- 50 pulses of 1.5 $\mu\text{s}$

Total ion-photon reaction rate = 250 Hz

### Momentum imaging from DET 1 and DET 2



### Momentum imaging from DET 1-3

Position and time

$$\rho_x = r_x/L_x, \quad \tau_x = t/(v_x L_x)$$

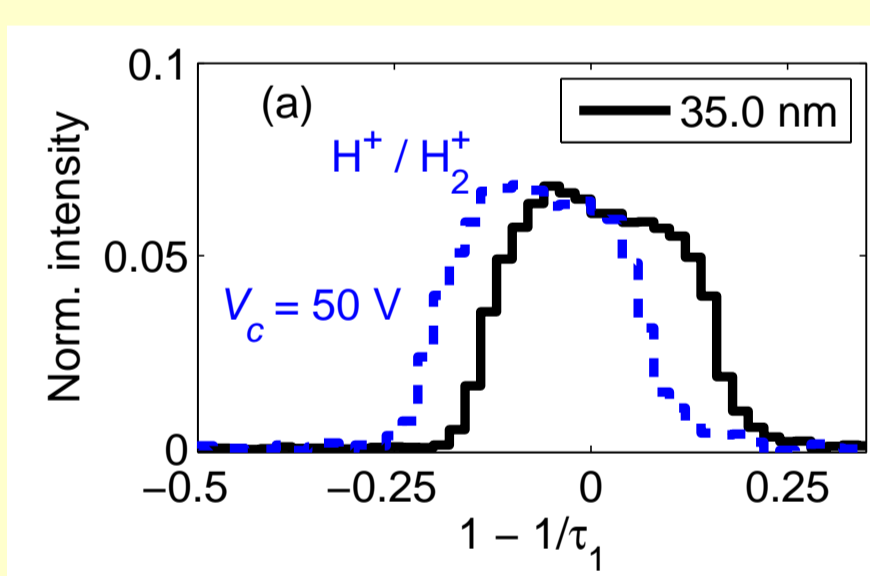
$$p_F = m_F v_F$$

$$\frac{\Delta p_x}{p_x} = \frac{1}{\tau_x} - 1$$

$$\frac{\Delta p_y}{p_y} = \frac{\rho_y}{\tau_y}$$

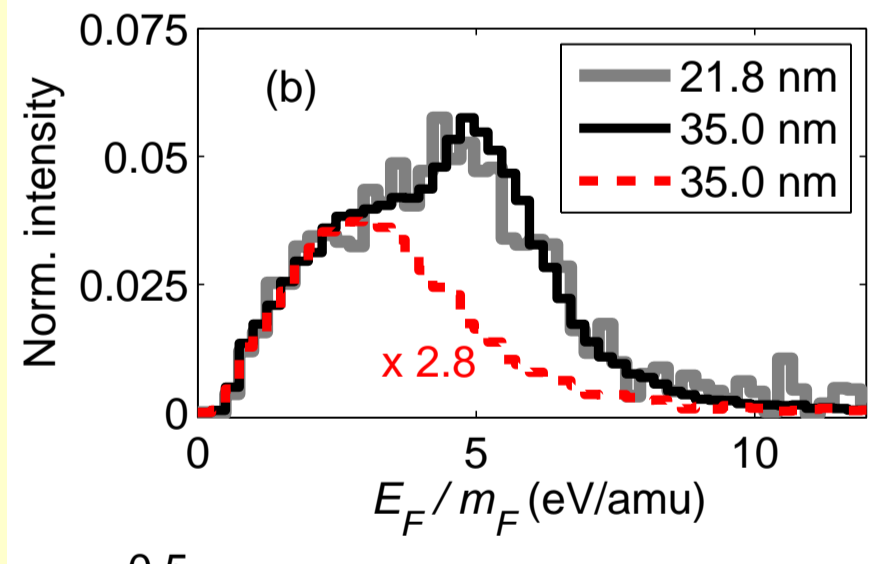
Relative momenta

## Fragmentation channel identification



### Relative longitudinal momentum on DET 1

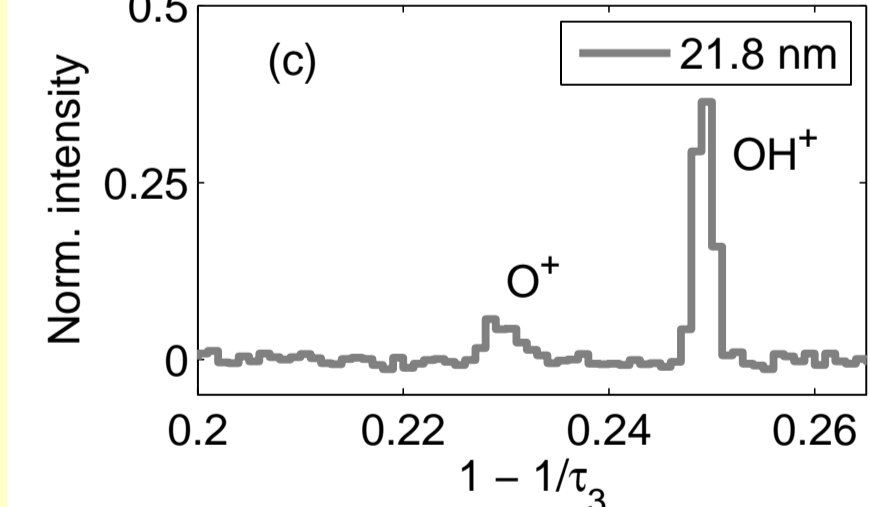
All light fragments are charged, i.e.  $H^+$  or  $H_2^+$



### Energy per mass on DET 1

Different wavelength - same break-up dynamics

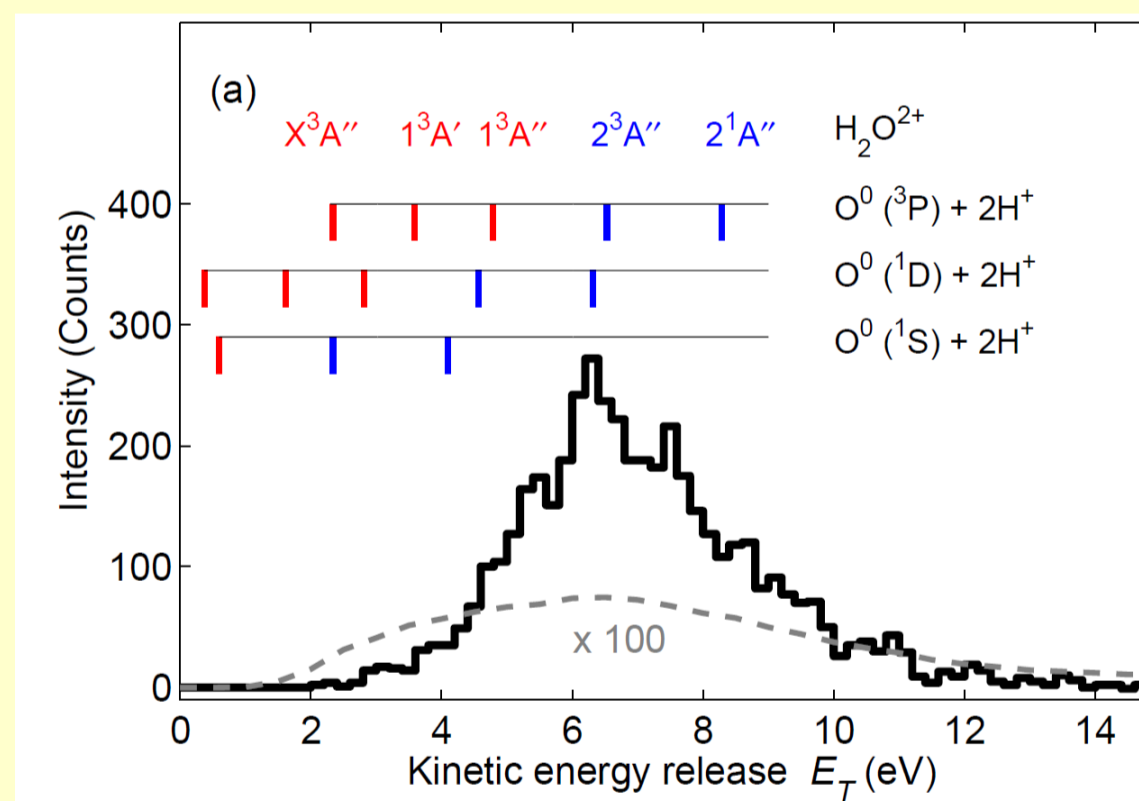
Coincidence with  $O^0$  (red curve)  
Three-body  $< 1.4$   
Two-body 1.0



### Relative longitudinal momentum on DET 3 (new)

Two-body channel  
 $OH^+ + H^+$  1.0  
 $O^+ + H_2^+$  0.2

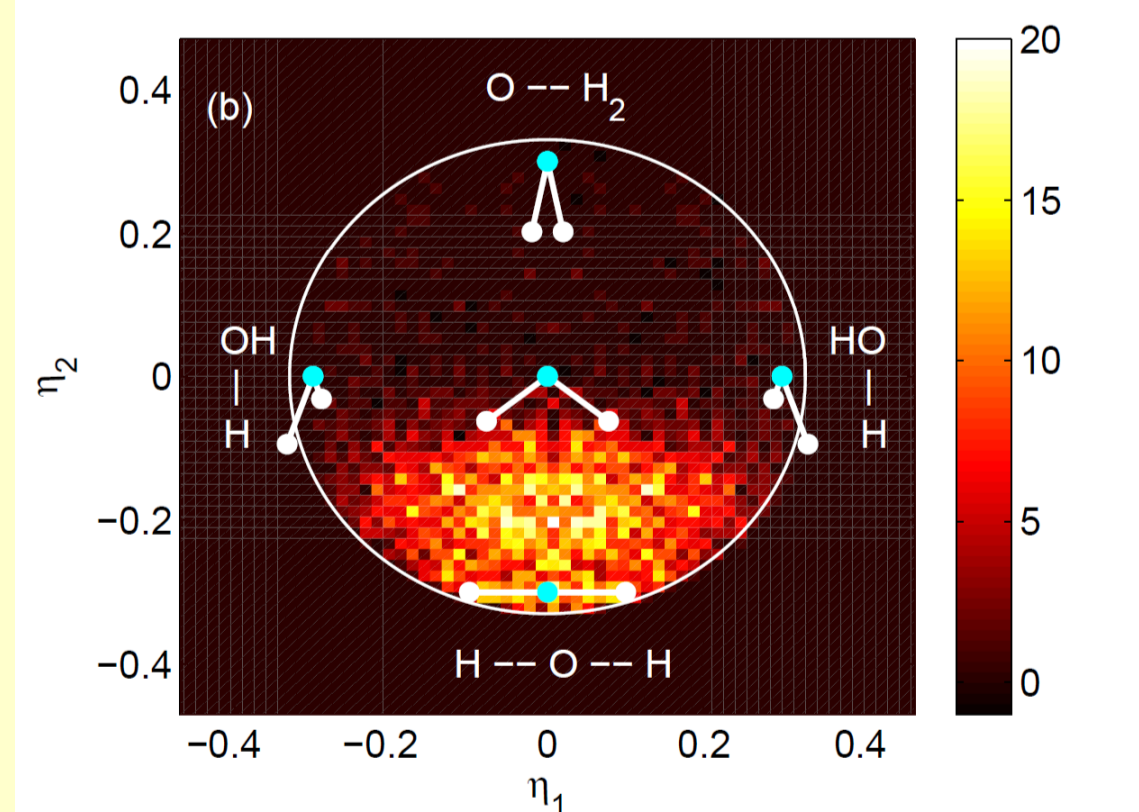
## $\gamma + H_2O^+ \rightarrow H_2O_2^{2+} \rightarrow O^0 + H^+ + H^+$



### Total kinetic energy release

$$E_T = \left[ \left( \frac{1}{\tau_i} - 1 \right)^2 + \left( \frac{\rho_i}{\tau_i} \right)^2 \right] \times \frac{m_F}{m_i} E_i$$

$$E_T = E_{O^0} + E_{H^+} + E_{H^+}$$

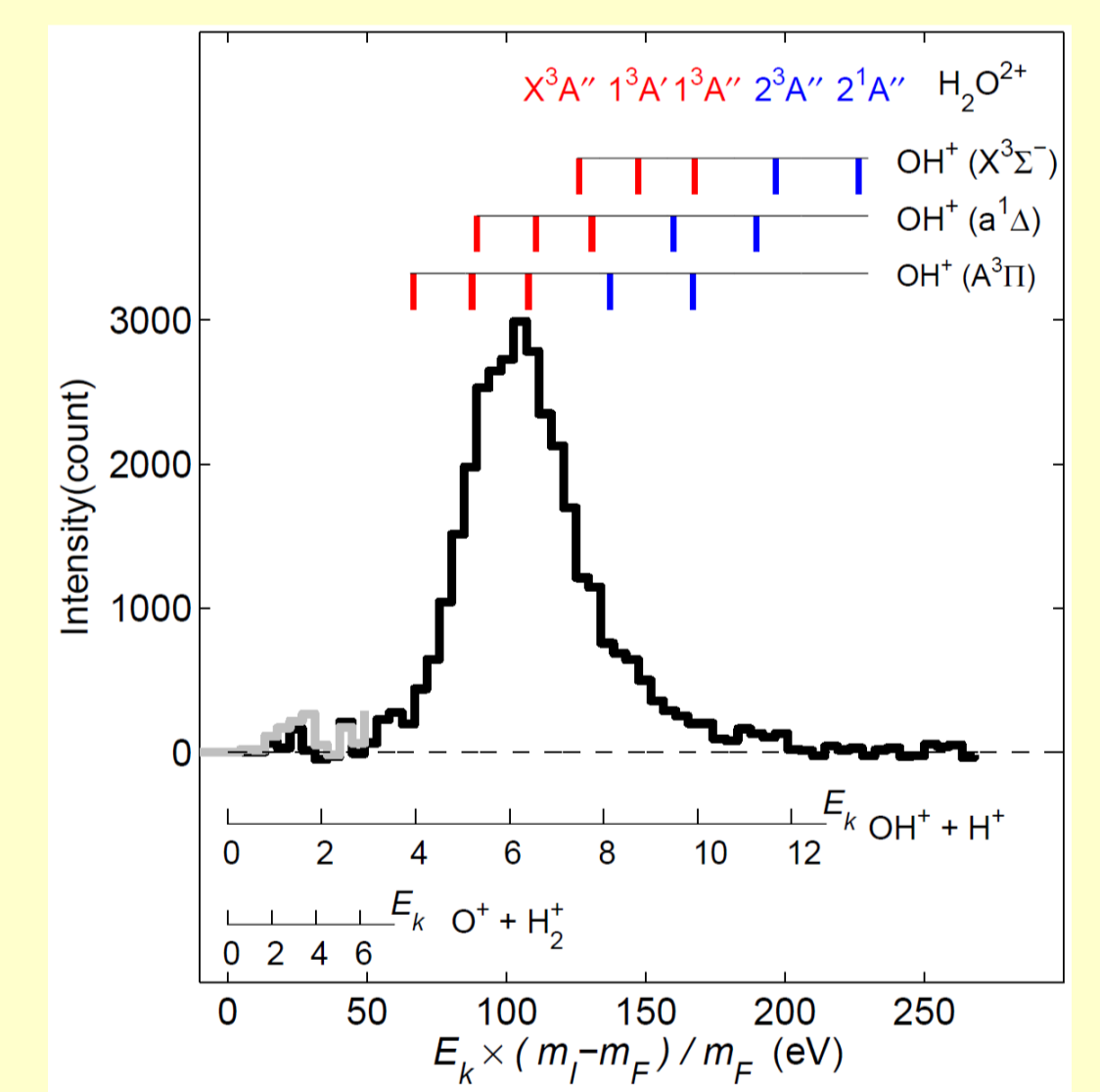


### 3-body break-up dynamics - Dalitz plot [17]

$$\eta_1 = \frac{m_i}{3m_o} \frac{E_{H^+} - E_{H^+}}{\sqrt{3E_T}}$$

$$\eta_2 = \frac{m_i}{3m_o} \frac{E_{O^0} - 1}{E_T - 3}$$

## $\gamma + H_2O^+ \rightarrow H_2O_2^{2+} \rightarrow OH^+ + H^+$



### Two-body kinetic energy release

$$E_k = \left[ \left( \frac{1}{\tau_i} - 1 \right)^2 + \left( \frac{\rho_i}{\tau_i} \right)^2 \right] \times \frac{m_F}{m_i - m_r} \times E_i$$

## Future directions with TIFF at FLASH

- For  $H_2O^+$** 
  - High statistics measurement with the new fragment analyzer to enable full characterization of the two-body channels
  - Photoelectron-photofragment spectroscopy to experimentally completely specify the photofragmentation pathways (photoelectronspectroscopy has been demonstrated for anions and larger water clusters)
- Larger molecular systems under study/proposal**
  - Water clusters ions  $H^+(H_2O)_n$  - Astrophysics / atmospheric physics
  - Small molecular cations:  $HeH^+$ ,  $He_2^+$ ,  $O_2^+$ ,  $N_2^+$ , ... - Non-adiabatic interactions / competition between ionization and excitation
  - Atomic anions  $He^-$ ,  $Li^-$ ,  $O^-$ , ... - full characterization of processes dominated by electron correlation
  - Molecular anions:  $H+(N_2)$ ,  $C_2H_2^+$ ,  $C_n^+$ ,  $PAH^+$ , ... - Astrophysics / atmospheric physics
- Collaborations are welcome!**

## Present conclusions on $H_2O^+$

- Photofragmentation dynamics similar at 35.0 nm and 21.8 nm. XUV induced process: ionization followed by fragmentation.
- Major fragmentation channels of  $H_2O^+$  under ionizing radiation
  - $O + H^+ + H^+$  1.0
  - $OH^+ + H^+$   $< 1.4$
  - $O^+ + H_2^+$  0.2
- Three-body-channel ( $O^0 + H^+ + H^+$ )
  - Wide distribution of energy release around 6 eV
  - Consistent with predictions of Gervais *et al.* [16]
  - Break-up dynamics: "the protons carries the momentum"
- Two-body-channel ( $OH^+ + H^+ / O^+ + H_2^+$ )
  - Major kinetic energy release around 5 eV
  - Future measurements can distinguish channels and clarify angular distributions
- Further interpretation of fragmentation mechanisms from theory would be very interesting!

## References

- [1] G. Herzberg, *Ann. Geophys.* **36**, 605 (1980).
- [2] P. Wehinger *et al.*, *Astronophys. J.* **190**, L43 (1974).
- [3] P. Wehinger and S. Wyckoff, *Astronophys. J.* **192**, L41 (1974).
- [4] A. H. Delsemme *et al.*, *Astronophys. J.* **209**, L153 (1976).
- [5] W.-H. Ip *et al.*, *Astronophys. J.* **293**, 609 (1985).
- [6] N. Meyer-Vernet *et al.*, *Astron. J.* **93**, 474 (1987).
- [7] M. A. Disanti, U. Fink, and A. B. Schultz, *Icarus* **86**, 152 (1990).
- [8] D. A. Neufeld *et al.*, *Astron. Astrophys.* **521**, L10 (2010).
- [9] H. Gupta *et al.*, *Astron. Astrophys.* **521**, L47 (2010).
- [10] V. Ossendorf *et al.*, *Astron. Astrophys.* **518**, L111 (2010).
- [11] M. Gerin *et al.*, *Astron. Astrophys.* **518**, L110 (2010).
- [12] E. F. van Dishoeck *et al.*, *Astrophys. J. Ser.* **62**, 109 (1986).
- [13] D. Smith, *Chem. Rev.* **92**, 1473 (1992).
- [14] D. J. Hollenbach *et al.*, *Annu. Rev. Astron. Astrophys.* **35**, 179 (1997).
- [15] D. J. Hollenbach *et al.*, *Astro. J.* **690**, 1497 (2009).
- [16] B. Gervais *et al.*, *J. Chem. Phys.*, **131**, 024302 (2009).
- [17] R. H. Dalitz, *Philosophical Magazine* **44**, 1068 (1953).